

## The LYRA Instrument Onboard PROBA2: Description and In-Flight Performance

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Received: ..... / Accepted: ..... / Published online: .....

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**Abstract** The *Large Yield Radiometer* (LYRA) is an XUV–EUV–MUV (soft X-ray to mid-ultraviolet) solar radiometer onboard the European Space Agency PROBA2 mission that was launched in November 2009. LYRA acquires solar irradiance measurements at a high cadence (nominally 20 Hz) in four broad spectral channels, from soft X-ray to MUV, that have been chosen for their relevance to solar physics, space weather and aeronomy. In this article, we briefly review the design of the instrument, give an overview of the data products distributed through the instrument website, and describe the way that data are calibrated. We also briefly present a summary of the main fields of research currently under investigation by the LYRAconsortium.

**Keywords:** Instrumentation and Data Management, Solar Irradiance, Flares, Earth’s atmosphere, Eclipse Observations

### 1. Introduction

The *Large Yield Radiometer* (LYRA: Hochedez *et al.*, 2006), is an XUV to MUV (soft X-ray to mid=ultraviolet) solar radiometer embarked on the European Space Agency PROBA2 mission that was launched on 2 November 2009. It revolves on a Sun-synchronous orbit with an altitude around 720 km, and

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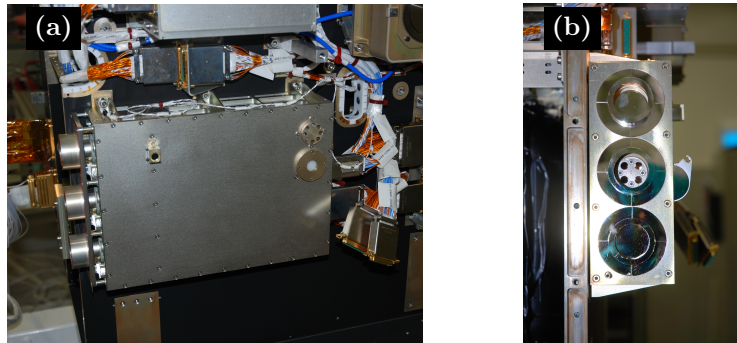
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faces the Sun continuously. LYRA acquires solar-irradiance measurements in four broad spectral channels, from soft X-ray to UV, that have been chosen for their relevance to solar physics, space weather and aeronomy.

The objective of the instrument is twofold:

- providing time series of solar irradiance with a very high sampling cadence (up to 100 Hz) in spectral ranges complementary to other active radiometers, *e.g.* *Geostationary Operations Environmental Satellite/Extreme Ultraviolet Sensor* (GOES/EUVS), *Solar and Heliospheric Observatory/Solar Extreme Ultraviolet Monitor* (SOHO/SEM: Judge *et al.*, 1998), *Solar Radiation and Climate Experiment/Solar Stellar Irradiance Comparison Experiment* (SORCE/SOLSTICE: Rottman, Woods, and Sparn, 1993; Woods, Rottman, and Ucker, 1993), *Thermosphere Ionosphere Mesosphere Energetics and Dynamics/Solar EUV Experiment* (TIMED/SEE: Woodraska, Woods, and Eparvier, 2004), *Picard/Precision Monitor Sensor* (PREMOS: Schmutz *et al.*, 2009), *Solar Dynamics Observatory/EUV Variability Experiment* (SDO/EVE: Woods *et al.*, 2012).
- testing an innovative kind of wide-bandgap diamond detectors. Such detectors, being radiation-hard and visible-blind, are particularly well adapted to observe the wavelengths targeted by LYRA.

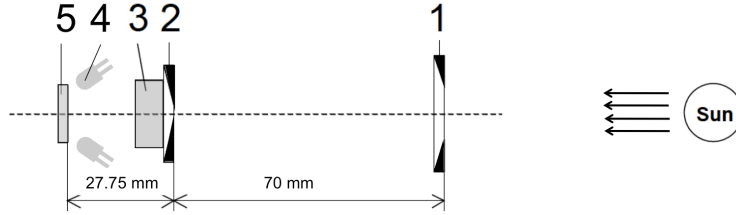
The instrument and its pre-flight calibration were described by BenMoussa *et al.* (2009). Here, we review the design of the instrument in Section 2. Section 3 provides an overview of the data products that are distributed to the scientific community and details their calibration. Section 4 briefly presents the main fields of scientific research currently investigated, exploiting LYRA data.



**Figure 1.** LYRA ( $315 \times 92.5 \times 222 \text{ mm}^3$ ) during its integration on PROBA2. Panel (b), the open cover of unit 2 reveals the four observation channels.

## 2. Design of the Instrument

LYRA is a shoe-box size instrument ( $315 \times 92.5 \times 222 \text{ mm}^3$ ) composed of three quasi-redundant units, each equipped with an individual cover and hosting four



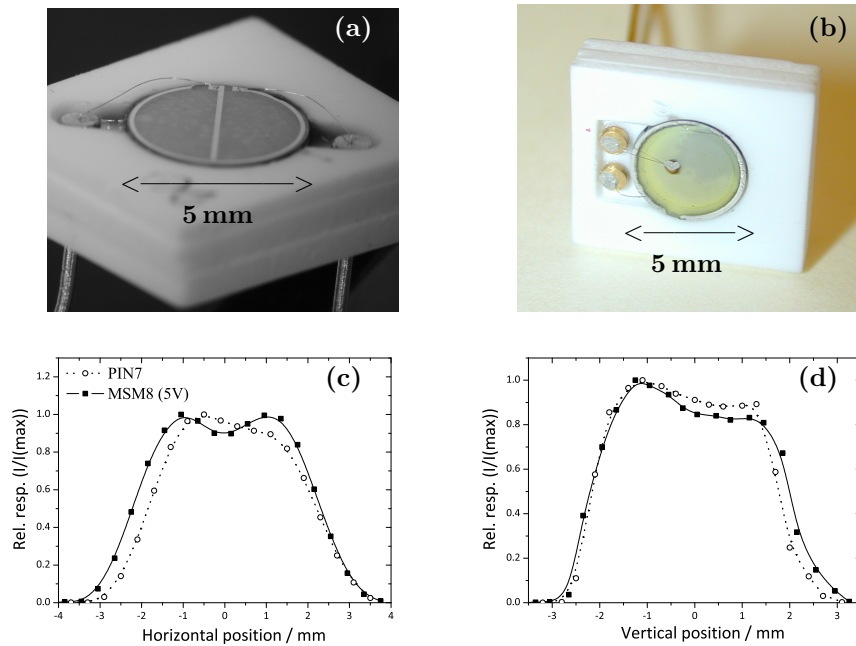
**Figure 2.** Schematic representation of one LYRA channel: 1 and 2 are respectively the view-limiting and precision apertures, together forming the collimator, 3 is the filter, 4 are the two LEDs at 375 and 470 nm located behind the filter, and 5 is the detector.

**Table 1.** Summary of the characteristics of LYRA channels. The purity is defined as the ratio of the flux in the nominal wavelength range (defined bandwidth) to the total output signal. Purities correspond to a solar-minimum-type spectrum. Thicknesses of aluminum and zirconium filter layers are indicated in the “filter label” column, together with the Acton reference for Lyman- $\alpha$  and Herzberg filters.

Channel	Filter label	Detector	Bandwidth	Purity
<i>Unit 1</i>				
1-1	Lyman- $\alpha$ [122XN]	MSM Diamond	120–123 nm	26 %
1-2	Herzberg [220B]	PIN Diamond	190–222 nm	95 %
1-3	Aluminum (158 nm)	MSM Diamond	17–80 nm + < 5 nm	96.8 %
1-4	Zirconium (300 nm)	AXUV Si	6–20 nm + < 2 nm	97 %
<i>Unit 2</i>				
2-1	Lyman- $\alpha$ [122XN]	MSM Diamond	120–123 nm	25.7 %
2-2	Herzberg [220B]	PIN Diamond	190–222 nm	95 %
2-3	Aluminum (158 nm)	MSM Diamond	17–80 nm + < 5 nm	97.2 %
2-4	Zirconium (141 nm)	MSM Diamond	6–20 nm + < 2 nm	92.2 %
<i>Unit 3</i>				
3-1	Lyman- $\alpha$ [122N+XN]	AXUV Si	120–123 nm	32.5 %
3-2	Herzberg [220B]	PIN Diamond	190–222 nm	95 %
3-3	Aluminum (158 nm)	AXUV Si	17–80 nm + < 5 nm	96.6 %
3-4	Zirconium (300 nm)	AXUV Si	6–20 nm + < 2 nm	95 %

spectral channels (see Figure 1). A channel consists of a collimator, an optical filter, a detector, and two LEDs on the side. LEDs are located between the filter and the detector and emit at 375 and 470 nm (see Figure 2). They are used to estimate the impact of ageing on the detectors.

The three units are essentially similar in terms of spectral coverage, but involve non-identical associations of filters and detectors (Table 1). The four bandpasses are indicated in this table, as well as their purity (*i.e.* the ratio of the flux in the defined bandwidth to the total output signal). Purities have been computed considering a quiet-Sun-type spectrum. They show that for all channels but Lyman- $\alpha$ , most of the detected signal actually comes from the defined bandpasses. The Lyman- $\alpha$  channel, however, is highly contaminated by out-of-band radiation.



**Figure 3.** Two examples of diamond detectors: MSM (a) and PIN (b). In both cases, we note the position of an electrode (central bar for the MSM, off-center circle for the PIN). Those electrodes strongly affect the flat-field of the detectors, as shown in the plots of relative responsivity along two perpendicular directions in (c) and (d).

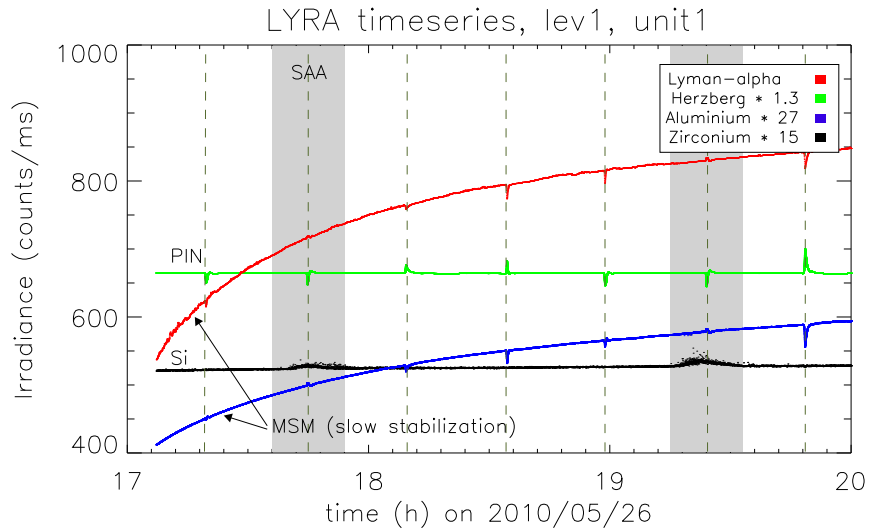
In addition to wide-bandgap diamond detectors prototypes, units 1 and 3 incorporate some classical silicon (Si) detectors to allow the comparison between both technologies. Diamond detectors are of two types: metal–semiconductor–metal (MSM) photoconductors and positive–intrinsic–negative (PIN) photodiodes (Figure 3), the latter being used for the Herzberg channels only (see Table 1). Characteristics of these detectors are detailed by BenMoussa *et al.* (2004) and BenMoussa *et al.* (2006). A few of these characteristics are reminded here, since they significantly affect LYRA data:

- An electrode is located on the surface of both MSM and PIN detectors, which results in important flat-field variations (see Figure 3).
- Trapping/detrapping of generated photoelectrons by defects (not bulk but surface defects) causes the signal of MSM detectors to take quite a long time before reaching stabilization (see Figure 4).

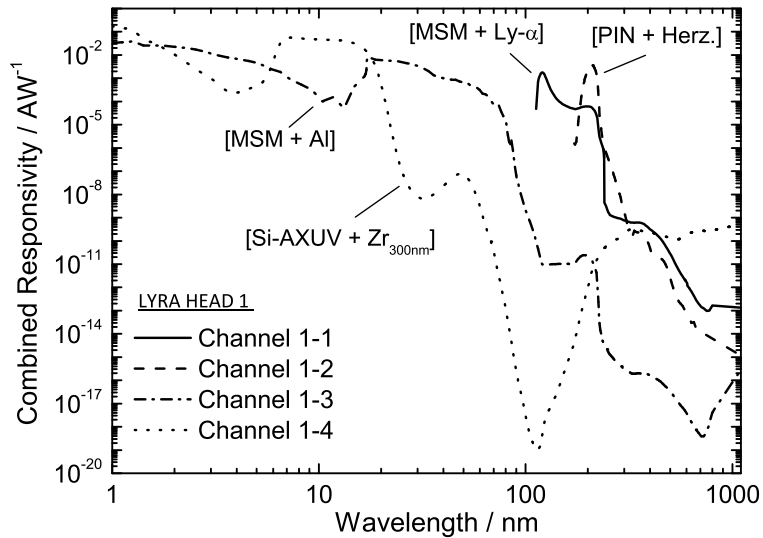
Figure 5 illustrates the total responsivity of filter-detector combinations for all four channels of unit 1. These curves result from a radiometric model which uses pre-launch measurements of filter transmittance and detector responsivity as input.

Raw LYRA data, as transmitted by the spacecraft, result from an on-board treatment of photocurrents produced by the detectors. Photocurrents are first





**Figure 4.** Slow stabilization of MSM detectors: the signal has not yet reached its stabilization level three hours after the detector was switched on. Also illustrated on this figure is the perturbation in the Si-detectors' signal when crossing the SAA. Diamond technology has proven to be more robust than silicon to the impact of high energy protons (and their secondary electrons) that causes this noisy behavior. Gray bands indicate the SAA and dashed vertical lines the large-angle rotations of the spacecraft. Time series have been rescaled to fit the same range and appear in the same order as in the legend. Scaling coefficients are indicated in the legend.



**Figure 5.** Simulated combined spectral (detector + filter) responsivities for LYRA unit 1 between 1 and 1000 nm, from BenMoussa *et al.* (2009).

converted to voltages in the 0–5 volts range by a resistor, of which the value is different for each channel. Voltages are then switched to frequencies by a voltage-to-frequency converter (VFC). Regularly, those VFC are connected to stable calibrated onboard voltage sources of 5 volts, 2.5 volts, and 0 volts. These measurements are interleaved in the telemetry, to be used as reference when processing the data. At the end of the acquisition chain there is a counter that counts the number of pulses received over the duration of the integration period. This number, which is expressed in counts, constitutes the raw LYRA data that are downloaded as part of the telemetry.

### 3. LYRA Data Description

#### 3.1. Data Products

LYRA produces time series of spectral irradiance in its four bandpasses in a quasi-uninterrupted way. Interruptions might happen during calibration campaigns (once every two weeks on average), or during the Winter occultation season, when the spacecraft transits the Earth’s shadow. Nevertheless, since PROBA2 is flying on a polar, dawn–dusk heliosynchronous orbit, such occultations are limited to a three–four months winter period (depending on the channel) and last for maximum 25 minutes every orbit (one orbit is 100 minutes long).

As mentioned in Section 2, LYRA includes three units, which are similar from the spectral point of view. The four channels of a unit are operated in parallel, acquiring irradiance measurements at a nominal cadence of 20 Herz, but that could go up to 100 Herz.

The strategy behind the redundancy is the following:

- The nominal unit (unit 2) is in permanent use, but it is therefore the most affected by degradation. In the Lyman- $\alpha$  and Herzberg channels of this unit, the signal has dropped by about 99% since the beginning of the mission. The degradation is so strong that the solar signal is now barely detectable in those two channels.
- Unit 3 is used in a campaign-driven way and keeps its cover closed the rest of the time, limiting ageing effects.
- Unit 1 is kept mostly unused, acquiring data for 40 minutes every three months on average and is therefore the most preserved. It is aimed at being a reference to estimate the degradation of units 2 and 3.

Data are usually available within four hours following their acquisition. Products with different levels of processing are distributed through the PROBA2 website (<http://proba2.oma.be>). They consist of uncalibrated and calibrated data, as well as various quicklook datasets (see Table 2).

Once downloaded, raw LYRA data in counts are converted back to counts  $\text{ms}^{-1}$ , dividing them by the integration time, and distributed without further processing through the instrument website as **level-1/engineering data**,

**Table 2.** Summary of LYRA data products distributed to the scientific community. These products are processed after each data download (*i.e.* every three to four hours).

Product	File extension on LYRA website	Format	Characteristics
Level 1 engineering data	*_lev1_std(bst).fits	FITS	unprocessed solar irradiance, in [counts ms <sup>-1</sup> ]
	*_lev1_cal(bca).fits	FITS	unprocessed calibration data, in [counts ms <sup>-1</sup> ]
	*_lev1_met.fits	FITS	ancillary data: temperature, pointing . . .
	*_lev1_rej(bre).fits	FITS	rejected samples (outliers . . .)
Level 2 basic science data	*_lev2_std.fits	FITS	calibrated solar irradiance, in [Wm <sup>-2</sup> ]
Level 3 averaged science data	*_lev3_std.fits	FITS	level 2 averaged over 1 min, in [Wm <sup>-2</sup> ]
Level 4 A quicklooks	*.png	image	daily plot of calibrated data for all LYRA channels
Level 4 B quicklooks	*.png	image	3-days GOES-like plot of calibrated data in aluminum and zirconium channels
Level 5 flare list	html	text file	List of flares with links to LYRA and GOES flux profiles

in standard files – either from the nominal unit (.std) or from an additional ‘back-up’ unit (.bst), that can be either unit 1 or 3. Information about the acquisition context (temperature, pointing, status of covers, status of LEDs, usage of a back-up unit . . .) compose the ancillaries stored in the metadata (.met) file. Additional files gather the data acquired during calibration campaigns (dark currents, LED signals), or the rejected data (outliers, data acquired during transitions between acquisition modes).

**Level-2/calibrated data** are available through the same website. Currently, calibration includes subtraction of dark currents (which removes the temperature effects), compensation for degradation, rescaling to one astronomical unit, and conversion from counts ms<sup>-1</sup> to irradiance units, but no correction for flat-field effects yet. Special features, such as imprints of large angle rotations of the spacecraft or perturbations due to the South Atlantic Anomaly, are also visible in the data. Those features are detailed in Section 3.3.

Basic IDL procedures to download and read LYRA data are available through the SolarSoft library ([sohowww.nascom.nasa.gov/solarsoft/](http://sohowww.nascom.nasa.gov/solarsoft/)) at the address [ssw/proba2/lyra/idl](http://ssw/proba2/lyra/idl).

## 3.2. Calibration of LYRA Data

### 3.2.1. Dark Current Subtraction

The dark currents were measured as a function of temperature between  $-40^{\circ}\text{C}$  and  $+60^{\circ}\text{C}$  in the laboratory before the launch, but only in steps of  $10^{\circ}\text{C}$ . The relationship between temperature and dark current was globally linear below  $40^{\circ}\text{C}$  (*i.e.* over the estimated operational temperature range), which explains why no further test was performed with smaller steps. Unfortunately, it turned out that the onboard temperature experienced in space was much hotter than expected – actually between  $+35^{\circ}\text{C}$  and  $+55^{\circ}\text{C}$  – in a range where the functional relationship between temperature and dark current is non-linear.

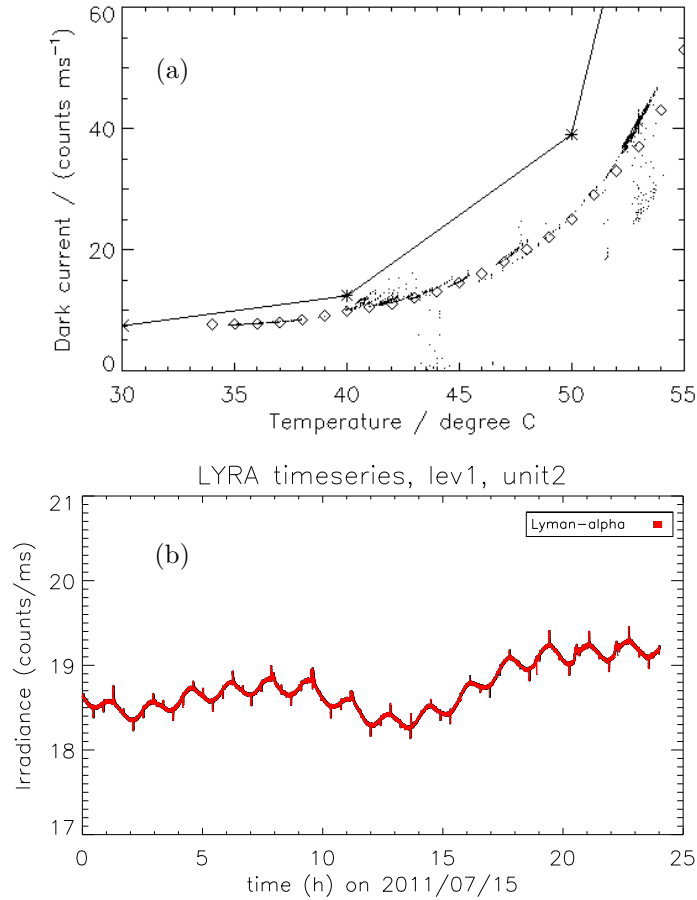
The actual relationship had therefore to be tabulated in smaller steps, exploiting several calibration campaign observations with closed covers, and even some with open covers, when the solar component could be removed. An example is illustrated in Figure 6 for the Lyman- $\alpha$  channel of unit 2 (one of the channels that is most affected by temperature effects). A similar relationship was found for each channel. The resulting dark-current estimation is finally subtracted from the data.

### 3.2.2. Correction for Degradation

From the very first hours after the covers were opened, LYRA suffered severe degradation, affecting more seriously its longer-wavelength channels (Figure 7). A possible explanation for the degradation observed is condensation and UV-induced polymerization of outgassing molecules on the filter surface. The resulting layer absorbs longer wavelengths more than shorter ones, explaining the different impact on different channels. In the Lyman- $\alpha$  and Herzberg channels of unit 2, the signal fell by 70 % within the first month alone and a loss of 99 % is now reached, with a degradation process tending to stabilize. The degradation in those channels is now such that solar signal is barely detectable in time series produced by the nominal unit. Fortunately, preserved back-up units are available for observation campaigns. It was not possible in the long run to calibrate and compare LYRA with other instruments without first taking the degradation into account. This, in turn, meant that the degradation had to be separated from the solar variation.

To some extent, the degradation can be calculated by internal means. For the two shorter-wavelength channels (aluminum and zirconium), this is done with the help of the spare units 1 and 3, which so far were only used for very occasional campaigns. Their covers were only opened for about 50 hours during the whole year 2010, while unit 2 was observing the Sun almost continuously since it saw the first light. The zirconium channels in units 1 and 3 have seen no apparent loss at all at the time of writing and any of them could be used as a reference for both the aluminum and zirconium channels of unit 2, after subtraction of the variations induced by solar activity. The evolution of the aluminum and zirconium channels is correlated.

Unfortunately, there is no such internal reference for the longer-wavelength channels (Lyman- $\alpha$  and Herzberg) of unit 2. Degradation was so rapid in those

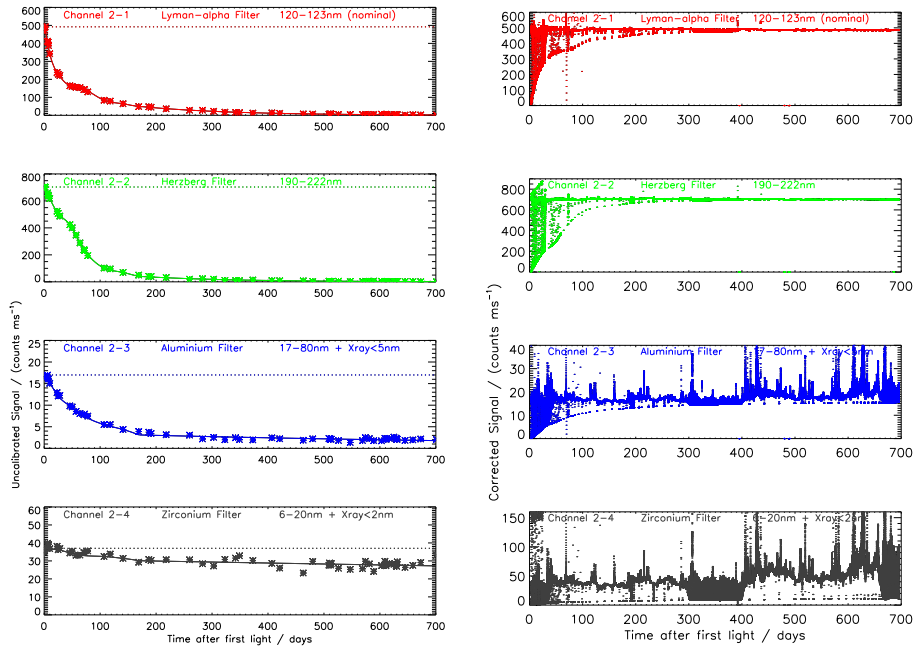


**Figure 6.** (a) Dark current in unit 2 Lyman- $\alpha$  channel *versus* temperature. Asterisks show pre-launch measurements, dots show data from calibration campaigns, diamonds show dark current final estimation. (b) An example of a quiet-day Lyman- $\alpha$  signal (level 1). The effect of orbital and daily temperature variations is clearly visible. After subtraction of dark current, the same time-series becomes a flat line.

channels that even units 1 and 3 were affected. The solution that we adopted is to conjecture that the quiet-Sun signal does not show any long-term trend. Any global deviation is considered as an effect of degradation. The data analysis is somewhat biased in this approach, and can only focus on short-term variations of the solar irradiance.

The evolution of the LYRA signal, especially in channels 2-1 and 2-2, shows phases of different degradation velocity:

- The degradation trend in the first half year (day 1–day 169 after first light) is fitted with a spline function through some manually selected data points. There is indeed no apparent mathematical function for the initial degradation; it seems to occur in various phases, and the physical processes



**Figure 7.** Left column: Degradation of unit 2 estimated from the beginning of the mission. The dark current is subtracted from the signal. Right column: Unit 2 data of the same period, after correction of degradation (and subtraction of dark current). During the season of occultations (days 0 to 40 and 300 to 400), the signal experiences drops to the dark-current level during every orbit. During the whole year, periodic calibration campaigns were scheduled, when the dark current signal was measured. As an unintended side effect of the additive correction for degradation, the level of the dark-current – which is also the level reached during occultations – gives the impression of a progressive increase with time, while being actually globally stable.

behind it (polymerization and emergence of contamination layers on the filters’ surfaces) are not well enough known.

- The degradation trend for the second half year was at first fitted with a function of the type  $1/(a+bt)$ , where  $t$  is time and  $a$  and  $b$  are fitting coefficients. This appeared to be a satisfactory first approach. Other functions, negative exponentials in particular, were also tested. The results of the latter proved to be an even more robust estimate for the future behaviour. A function of the type  $\exp(a+bt)$  has been introduced in a recent update of the calibration software.

After day 100, the degradation has significantly slowed down, however. Therefore, only later data points, representing calibration campaigns after 23 June 2010, are used to estimate the future development.

The loss caused by degradation is estimated relative to the first light individually for each channel. The degradation is corrected by adding this loss to the current levels. Correction by addition (as opposed to multiplication by a correction coefficient) has advantages, because degradation is a function of the wavelength and does not uniformly affect the broad LYRA channels over

their whole spectral range. In particular, there is no apparent sensitivity loss in the X-ray range. Flares of similar intensity in GOES/XRS scale now peak to the same count rates in aluminum and zirconium channels as they did at the beginning of the mission, while the measurements of the overall EUV background in the same channels has significantly decreased. A multiplicative correction coefficient would artificially exaggerate the flares. On the other hand, there are some disadvantages: the variation of EUV components in channel 3 (and to a lesser degree in channel 4) might be underestimated, and the occultation profiles become distorted since they no longer drop to zero. These effects can be seen in the right panel of Figure 7.

### 3.2.3. Conversion into Physical Units

Photocurrents measured by LYRA detectors can be modeled by

$$i = i_s + i_d = \frac{A}{T} \int_t \int_\lambda E(\lambda, t) F(\lambda) D(\lambda) d\lambda dt + i_d \quad (1)$$

where :

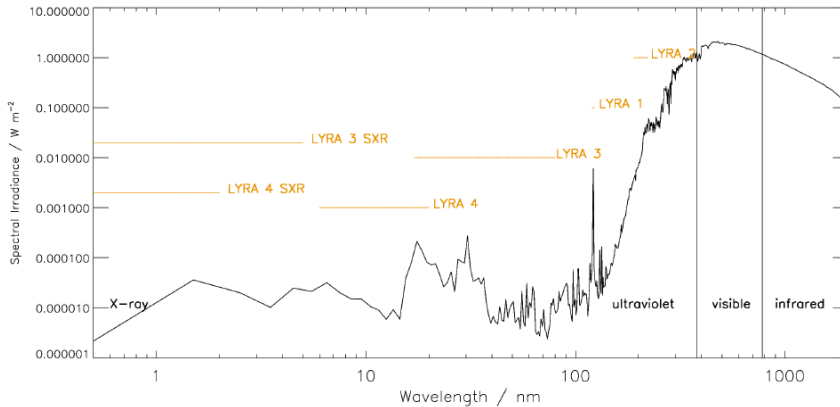
- $i$  is the measured photocurrent that can be defined as the sum of a solar [ $i_s$ ] and a dark current [ $i_d$ ] contribution
- $\lambda$  is the wavelength
- $t$  is the time and is integrated over an exposure
- $A$  is the aperture area, *i.e.* the area of the detector that is exposed
- $T$  is the total exposure time (nominally 50 ms)
- $E(\lambda, t)$  is the solar spectral irradiance
- $F(\lambda)$  is the filter transmittance
- $D(\lambda)$  is the detector spectral responsivity

and where the integral over  $\lambda$  is performed over the whole spectral range in which the instrument is sensitive (not restricted to the defined bandpass), which means that it also includes the out-of-band radiation. This out-of-band radiation constitutes a source of measurement error that might be important in the case of the Lyman- $\alpha$  channel (see purities in Table 1).

Since LYRA channels cover broad spectral ranges, it is not possible to directly invert Equation (1) to retrieve the spectral irradiance [ $E(\lambda, t)$ ] from the measured photocurrents [ $i$ ]. For an absolute radiometric calibration, we compare data acquired at any time [ $t$ ] to a pre-degradation reference LYRA measurement (the first light data of 6 January 2010), for which this conversion into irradiance units is known. This comparison is performed after correction for the degradation, as detailed in Section 3.2.2. It assumes that the relationship between LYRA count rate and irradiance in physical units is linear.

$$E_{\text{cal}} = \frac{i_{\text{uncal}} - i_d + \text{corr}}{i_{\text{uncal}}^{\text{FL}} - i_d^{\text{FL}}} E_{\text{cal}}^{\text{FL}} \quad (2)$$

where:



**Figure 8.** Reference solar spectrum reconstructed from TIMED/SEE and SORCE/SOLSTICE level-3 data on 6 January 2010, used as a reference to infer LYRA radiometric calibration. Wavelength coverage of different LYRA channels is marked with orange lines.

- $E_{\text{cal}}$  and  $E_{\text{cal}}^{\text{FL}}$  are the spectral irradiances covering one LYRA channel [ $\text{W m}^{-2}$ ] for respectively the current measurement and the first-light reference
- $i_{\text{uncal}}$  and  $i_{\text{uncal}}^{\text{FL}}$  are the unprocessed solar irradiance [ $\text{counts ms}^{-1}$ ] for respectively the current measurement and the first light reference
- $i_{\text{d}}$  and  $i_{\text{d}}^{\text{FL}}$  are the dark current measurements [ $\text{counts ms}^{-1}$ ] for respectively the current measurement and the first-light reference
- $\text{corr}$  is the corrective term for degradation

The conversion of this LYRA reference into irradiance units can be done by comparing it to measurements provided by other instruments.

Expressing the first-light reference measurement as physical irradiance units requires knowledge of the detailed solar spectrum for this day. We used a concatenation of spectral irradiance by TIMED/SEE from 0.5 nm to 115.5 nm and by SORCE/SOLSTICE from 116.5 nm to 2412.3 nm (see Figure 8).

This spectrum was then inserted in Equation (1) to produce an estimation of the expected LYRA photocurrent. The difference of the estimated photocurrent to the measured one was converted into an excess/default of spectral irradiance with respect to the one observed by SEE and SOLSTICE (integrated over the LYRA spectral range).

In other words, we calculated the LYRA first-light spectral irradiance in a given channel using

$$E_{\text{cal}}^{\text{FL}} = \frac{i_{\text{uncal}}^{\text{FL}} - i_{\text{d}}^{\text{FL}}}{\frac{A}{T} \int_t \int_{\text{sol.spec.}} E_{\text{S}}(\lambda, t) F(\lambda) D(\lambda) d\lambda dt} \int_{\text{bandpass}} E_{\text{S}}(\lambda) d\lambda \quad (3)$$

where  $E_{\text{S}}$  is the solar spectral irradiance (full spectral resolution) from SEE-SOLSTICE on 6 January 2010.

In Equation (3), the integrals over  $\lambda$  cover either the whole solar spectrum (denominator) or only the official spectral range of a channel (numerator), since the latter estimates what would have been measured by a perfect instrument.



**Table 3.** Excess of spectral irradiance as observed by LYRA in comparison to TIMED/SEE and SORCE/SOLSTICE measurements.

Lyman- $\alpha$	Herzberg	Aluminum	Zirconium
?	+18.0 %	+13.3 %	+9.2 %

An additional complication came from the fact that, in the case of the Lyman- $\alpha$  channel, the excess values computed for the three redundant units differed significantly (unit 1: +81.3 %, unit 2: +91.2 %, unit 3: +3.3 %). Moreover, it is difficult to compare units 1 and 2 (MSM diamond detectors) to unit 3 (Si detector), since the diamond detectors have an additional second peak around 200 nm, while the silicon detector collects 70 % of its non-nominal input between 200 and 1100 nm, with a peak between 900 and 1000 nm. Therefore, it is hard to make a statement such as “LYRA observes  $x$  % more irradiance as compared to SORCE/SOLSTICE” – which explains the question mark in Table 3. For this channel, the SORCE/SOLSTICE value for the nominal interval 120–123 nm must be assumed.

The other channels showed more consistency, so we used the average value over the three units (see Table 3)

Combining Equations (2) and (3), and taking into account that the solar irradiance might be considered as constant over sub-second periods, we obtain:

$$E_{\text{cal}} = \frac{i_{\text{uncal}} - i_{\text{d}} + \text{corr}}{A \int_{\text{sol.spec.}} E_{\text{S}}(\lambda) F(\lambda) D(\lambda) d\lambda} \int_{\text{bandpass}} E_{\text{S}}(\lambda) d\lambda \quad (4)$$

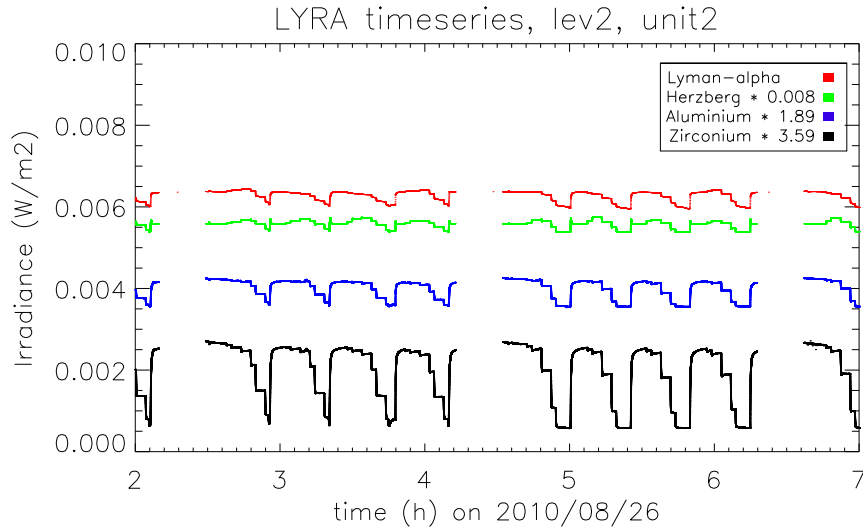
### 3.3. Non-Solar Features in the LYRA Data

Undesired features, such as imprints of large angle rotations of the spacecraft or perturbations due to the South Atlantic Anomaly (SAA), are present in LYRA data and must not be interpreted as solar signal variation. This section provides a list of those features.

#### 3.3.1. Flat-Field Effects

As described in Section 3.2, the calibration process does not yet correct the data for variations associated with pointing fluctuations. An analysis of PROBA2 attitude over several orbits reveals that pointing is stable up to 90 arcseconds. Spacecraft jitter introduces fluctuations in the LYRA signal of less than 1 %.

Nevertheless, it often happens that PROBA2 is off-pointed in the framework of calibration or scientific campaigns, introducing signal fluctuations of which the amplitude depends on the new pointing (see Figure 9). Such campaigns are noted in the public calendar of LYRA activities linked from the website of the instrument <http://proba2.oma.be/about/operations>. It is intended to correct for those fluctuations in a future version of the calibration routine.



**Figure 9.** Flat-field campaign of 26 August 2010: the spacecraft is off-pointed from  $0^\circ$  to  $3^\circ$ , in steps of  $0.5^\circ$ , successively in the S, E, N, W, SE, NE, NW, and SW directions. Results are plotted for unit 2 channels. Time series have been rescaled to fit the same range and appear in the same order as in the legend. Scaling coefficients are indicated in the legend.

### 3.3.2. Large-Angle Rotations of the Spacecraft

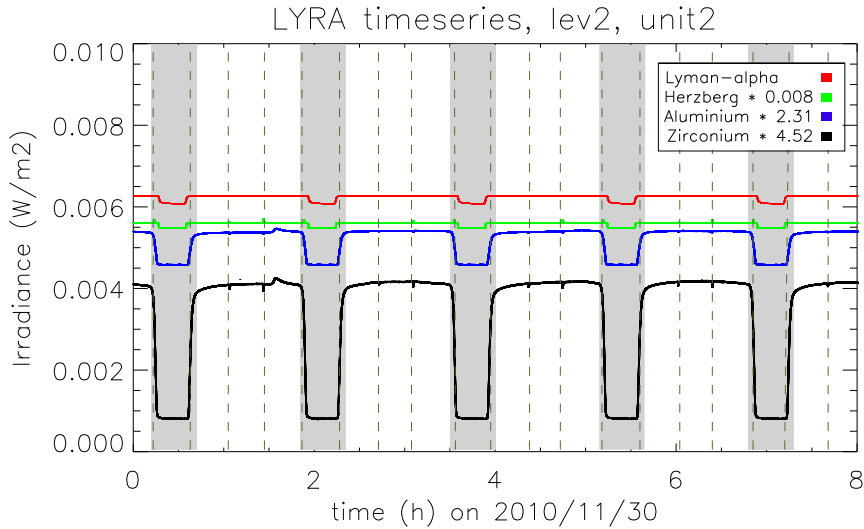
Four times per orbit, the spacecraft rotates  $90^\circ$  around the axis pointing toward the Sun, to avoid the Earth shadowing its star trackers. Because of the inhomogeneous flat field of the diamond detectors, such rotations are clearly visible in the time series (see Figures 10 to 12). Unfortunately, the acquisition of pointing parameters is performed at a limited cadence, not fast enough to deduce the spacecraft movement during such maneuvers with a sufficient accuracy and to allow for any flat-field correction. Large-angle rotations are systematic and cannot therefore be confused with natural solar variability. In the future, we are considering removing those features from the data.

### 3.3.3. Occultations

From November to February (approximately), the orbit of PROBA2 crosses the Earth's shadow. This produces progressive attenuation of the solar signal when LYRA is observing the Sun through deeper layers of the Earth's atmosphere and ends up with a total extinction, see Figure 10. As for large-angle rotations, those features are easily identifiable by an observer because of their regularity.

### 3.3.4. Slow Stabilization of MSM Detectors

MSM detectors need time to stabilize when exposed to light (see, *e.g.*, Figure 4). We can explain this phenomenon by the existence of surface defects, which trap the photoelectrons produced, preventing them from being collected at the



**Figure 10.** Drop of signal when the spacecraft transits the Earth's shadow (highlighted with gray). Vertical dashed lines indicate the large-angle rotations of the spacecraft. Time series have been rescaled to fit the same range and appear in the same order as in the legend. Scaling coefficients are indicated in the legend.

electrodes. The same happens when closing the covers: the MSM detector signal does not drop immediately to zero. Trapped electrons may even take several hours to leak out.

### 3.3.5. Perturbations in the LYRA Electronics

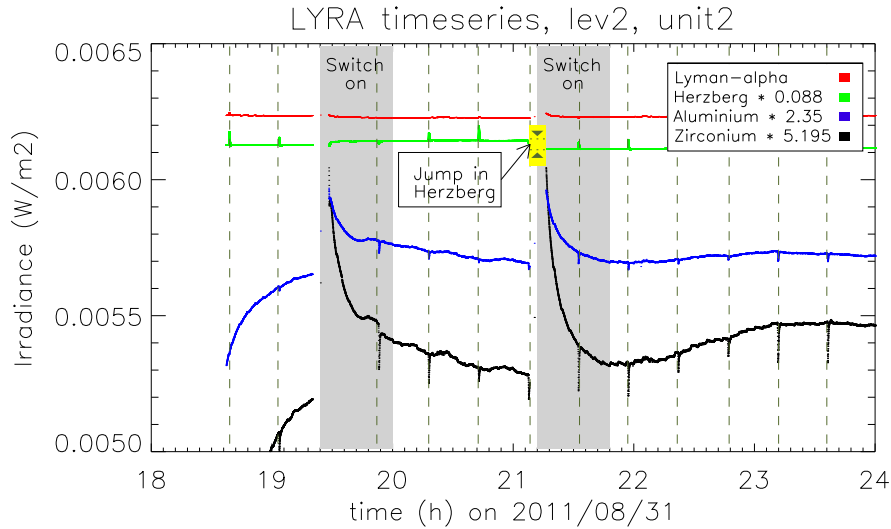
Switching on the instrument can produce a peak in the signal measured in some channels. This effect is likely to happen any time that we change the operational mode, *i.e.* when activating a new unit, performing a calibration campaign, reloading the onboard Field-Programmable Gate Arrays (FPGA) or coming back to nominal acquisition after one of the aforementioned activities. Such peaks are illustrated in Figure 11.

### 3.3.6. Jumps of Irradiance Observed in the Herzberg Channel

From time to time, the signal measured in the nominal Herzberg channel jumps from one level to another (small offset added or removed, see Figure 11). The reason for these jumps is not completely understood at the moment, but it is synchronous with the release of onboard commands. Jumps definitely complicate the exploitation of nominal Herzberg time series, whenever a high accuracy is required for solar irradiance.

### 3.3.7. South Atlantic Anomaly Perturbations

When transiting the South Atlantic Anomaly (SAA), secondary electrons generated by high-energy protons hit the detectors and make LYRA signal more noisy.

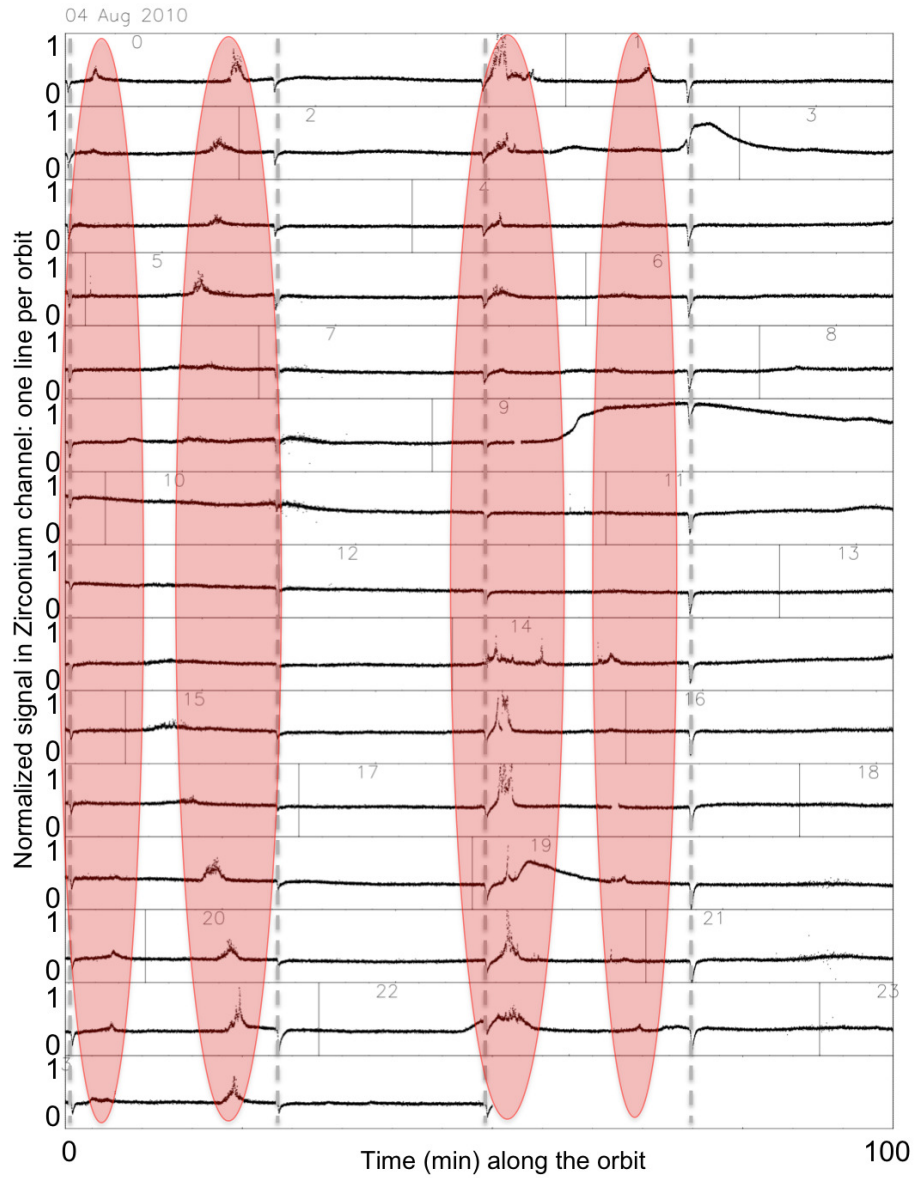


**Figure 11.** Typical perturbation in instrument electronics. The peaks in irradiance measurements in the aluminum and zirconium channels (highlighted with gray) are caused by switching the instrument on. A small jump in Herzberg channel irradiance is also visible. The occurrence of such offsets is synchronous with a change of acquisition parameters (*e.g.* change of the unit used, of the acquisition cadence or of the cover status). Dashed vertical lines indicate the large-angle rotations of the spacecraft. Time series have been rescaled to fit the same range and appear in the same order as in the legend. Scaling coefficients are indicated in the legend.

This effect mostly affects Si detectors, independently of the spectral range, while MSM and PIN detectors – being radiation hard – usually do not show significant perturbations (the effect of SAA on LYRA time series is visible in Figure 4). It is worth mentioning here that four channels of LYRA are amplified tenfold on-board to make up for their otherwise low signal: all three Lyman- $\alpha$  channels and the zirconium channel of unit 2. In these channels, the SAA perturbations appear magnified.

### 3.3.8. Auroral Perturbations

Auroral zones usually do not have any impact on the LYRA signal, with one noticeable exception: such perturbations appear in case of geomagnetic-storm conditions, usually when the Kp index is above four (see Figure 12). We have not yet identified what mechanism provokes the perturbations that we observe in LYRA data. Are they caused by auroral particles, subsequent bremsstrahlung, or auroral photons? The question is still under investigation. It is quite remarkable that it affects equally diamond and silicon detectors, but not equally all spectral ranges: aluminum and zirconium channels are affected, Lyman- $\alpha$  and Herzberg are not. On the other hand, the imager SWAP onboard PROBA2 (Seaton *et al.*, 2011), although observing at 17.4 nm, does not exhibit this effect; this may be related to the fact that LYRA, unlike SWAP, has a direct optical path, without any reflection by mirrors. In SWAP, the mirrors inserted in the optical path might absorb energetic particles, preventing them to reach the detector.



**Figure 12.** Perturbations appearing when crossing the auroral zone during or after a geomagnetic storm. Each horizontal line corresponds to one orbit. Four zones of perturbations are identified for each orbit (highlighted with red). They occur systematically in the same range of latitudes (north and south auroral ovals are each crossed twice during an orbit). Dashed vertical lines indicate the large-angle rotations of the spacecraft.

### 3.4. Radiometric Accuracy

From Equation (4), the maximal uncertainty on the calibrated LYRA data is described by

$$\frac{\Delta E_{\text{cal}}}{E_{\text{cal}}} = \frac{\Delta i}{i} + \frac{\Delta \left( \int_{\lambda} E d\lambda \right)}{\left( \int_{\lambda} E d\lambda \right)} + \frac{\Delta A}{A} + \frac{\Delta \left( \int_{\lambda} EFD d\lambda \right)}{\left( \int_{\lambda} EFD d\lambda \right)} \quad (5)$$

where  $i = i_{\text{uncal}} + i_{\text{d}} + \text{corr.}$  We will discuss the terms of this equation one by one.

#### 3.4.1. Error on the Measured Currents

The first term of Equation (5) takes into account the measurement error on the solar signal and on the dark current, as well as an estimation of the error introduced by the corrective term. Both solar signal and dark current depend on:

- a quantization error
- the stability of three onboard reference voltages (0 V, 2.5 V and 5 V), which are used to convert data expressed in count number to voltages. This conversion implies a polynomial fit (second order) of the three reference voltages, which also introduces its own error
- the uncertainty on the internal resistor, used to convert the voltages into currents
- the uncertainty on the integration time, which is related to the quartz stability

BenMoussa *et al.* (2009) estimated the relative uncertainty associated with these four parameters for each LYRA channel, and obtained an associated error of 0.03 % at the maximum.

Additionally,  $i_{\text{uncal}}$  is also affected by the pointing stability (jitter). For Sun-centered acquisitions, the jitter combined with flat-field non-homogeneities result in an uncertainty  $\leq 1\%$  in all channels.

Furthermore,  $i_{\text{uncal}}$  can be split into its in-band and out-of-band components. The latter is considered as part of the measurement error and is of the same order as the complement to the purity in Table 1. This is one of the major sources of uncertainty since it is of the order of 5 % of  $i_{\text{uncal}}$  in most channels and even of 75 % in Lyman- $\alpha$ .

Last but not least, the error associated to the corrective term is hard to estimate. But this term is based, for the aluminum and zirconium channels, on the assumption that the unit 3 zirconium channel did not degrade, and for Lyman- $\alpha$  and Herzberg channels on the assumption that they do not show any long-term trend. Therefore, one might consider that the error introduced by the corrective term is of the order of the actual degradation of the unit 3 zirconium channel (a few per cent of the corrective term) for aluminum and zirconium channels, and of the maximum variation over the solar cycle for Lyman- $\alpha$  (between 40 and 100 % of the corrective term) and Herzberg (about 5 %).

In conclusion, the uncertainty on the measured current is mostly due to the correction for degradation and to the out-of-band radiation, with the importance of the first one progressively increasing over the mission, as the second one decreases. On average, we consider an overall uncertainty on the measured current of about 100 % in Lyman- $\alpha$  and of 6 % in the other channels.

#### 3.4.2. Error on the Spectrum

The second term in Equation (5) depends on the accuracy of SORCE/SOLSTICE and TIMED/SEE measurements, which are of the order of 5 % and 10–20 % respectively. The first value is applicable to Lyman- $\alpha$  and Herzberg channels, while the second one is used for aluminum and zirconium channels.

#### 3.4.3. Error on the Aperture Area

The aperture area was measured at the Swiss Federal Office of Metrology and Accreditation (METAS), with an uncertainty of 0.07 % (see BenMoussa *et al.* (2009) ).

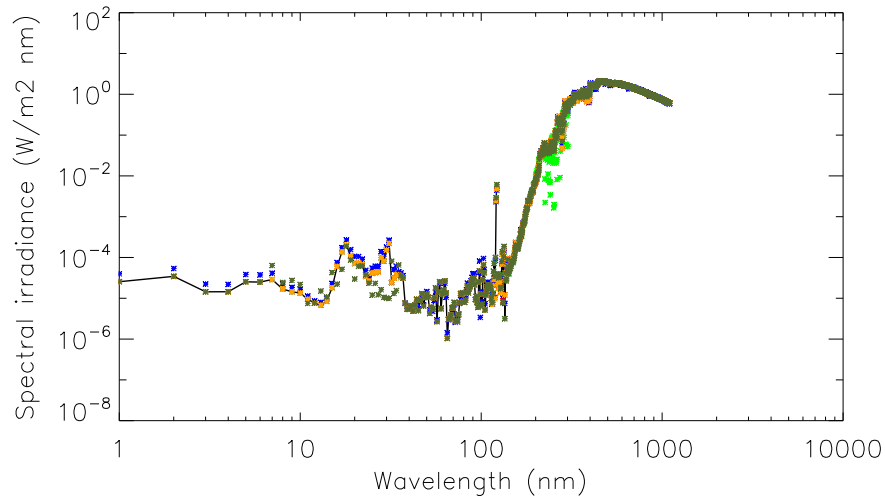
#### 3.4.4. Error on the Simulated Current

Combining the uncertainty on the filter and detector characterization, which is provided in BenMoussa *et al.* (2009), with that of the solar spectrum results in an estimated error of 15–20 % in all channels

#### 3.4.5. Conclusion

All in all, we consider the uncertainty on LYRA calibrated data to be of the order of 30–40 % for all channels except Lyman- $\alpha$ , where it is about 120 %.

In an attempt to validate those uncertainties, we have used alternative spectra in Equation (4) and have seen how LYRA data were affected. Unfortunately, for the considered date (6 January 2010) and spectral ranges, we are not aware of any other measured spectrum than the TIMED/SEE and SORCE/SOLSTICE ones. Even empirical models such as NRLSSI (Lean *et al.*, 2005) and SATIRE (Krivova *et al.*, 2009) are not provided for periods after 2006 and 2007 respectively. We have therefore tried to find a date with a similar solar-activity context as during LYRA first-light, for which those models produced spectra. We chose 10 November 2005. We also tested the SRPM model (Fontenla *et al.*, 2009), which applies for very quiet-Sun conditions. Additionally, to also cover the shortest wavelength ranges in LYRA, we picked a date after 6 January 2010 when SDO/EVE was in use and repeated the exercise. The selected day is 16 June 2010. The obtained spectra are plotted in Figure 13. While NRLSSI affected the four channels of LYRA, SATIRE and SRPM only concerned the wavelengths above 115 and 200 nm and therefore had a limited impact on the aluminum and zirconium channels. SDO/EVE (6–36 nm) on the other hand only had an impact on these two channels. The variations induced in LYRA measurements are summarized in Table 4 and are all within the cumulative error related to spectrum and simulated current.



**Figure 13.** Alternative spectra used to calibrate LYRA data. The full black line represents the TIMED/SEE-SORCE/SOLSTICE reference spectrum. Alternative spectra that were tested are over-plotted and respectively correspond to NRLSSI (blue), SRPM (green), SATIRE (orange), and EVE (olive).

**Table 4.** Variation percentage in LYRA calibrated data induced by the use of an alternative reference spectrum (these values are for LYRA nominal unit).

Spectrum	Lyman- $\alpha$	Herzberg	Aluminum	Zirconium
NRLSSI	-17.5 %	-0.4 %	-7.8 %	+0.6 %
SATIRE	+19.1 %	-3.5 %	0 %	0 %
SRPM	+2.3 %	-4.2 %	0 %	0 %
EVE	-	-	+3.7 %	+24.1 %

#### 4. Science Opportunities

In this section, we briefly describe the main topics of the scientific exploitation of the LYRA data.

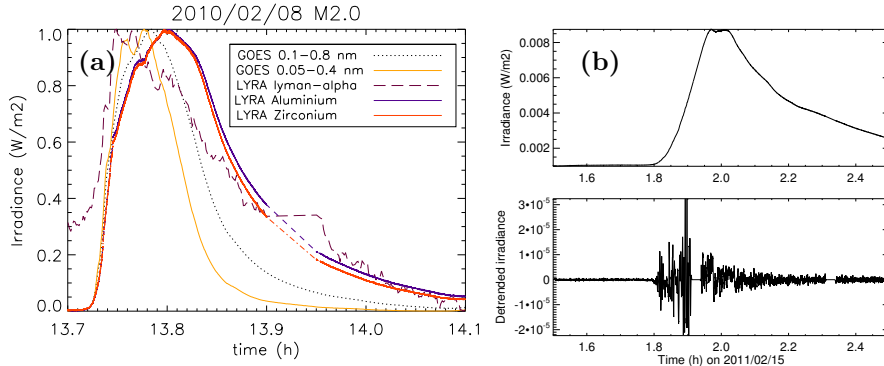
##### 4.1. Flares

Flares are one of the main scientific targets of LYRA. Comparing the flare profiles in LYRA channels with observations of instruments acquiring at other wavelengths (GOES/XRS, SDO/EVE...) allows one to determine a chronology of the temperature evolution in the flare and to confront it with the theoretical scenarios for energy release (see Figure 14, panel (a)).

In addition, the LYRA high sampling rate allows one to detect short-timescale phenomena, such as quasi-periodic pulsations (QPP) (Figure 14, panel (b)) – with periods from fractions of a second to a few minutes – that are observed



in flares in many wavelengths (see Van Doorselaere *et al.*, 2011; Dolla *et al.*, 2012). Here again, LYRA can help with completing the picture when trying to understand the origin of these QPP.



**Figure 14.** Flare analysis with LYRA data: (a) shows the differences between start-, end-, and peak-time of a flare when looking at various spectral ranges, from soft X-rays (GOES/XRS) to EUV (LYRA). For clarity, data corresponding to a large angle rotation of the PROBA2 spacecraft between 13:54 and 13:57 have been removed. (b) highlights oscillations detected during the onset of the X2.2 flare of 15 February 2011. The upper panel shows the flare curve in the LYRA zirconium time series, while the lower panel shows the same observations detrended by subtracting the signal smoothed using a 20-second boxcar.

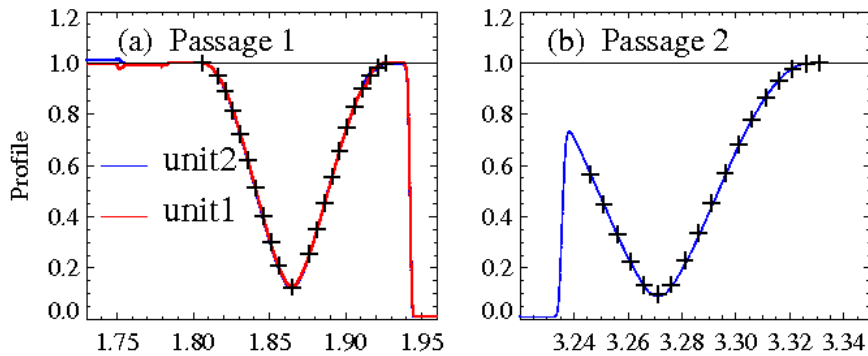
#### 4.2. Variability of Solar Irradiance

The Sun’s primary influence on Earth is through its radiation, and monitoring and understanding its variations are of prime importance. It is now well accepted that monitoring the solar spectral irradiance over the long-term is necessary in order to understand the role of the Sun in Earth’s climate change, especially the EUV part of solar spectrum, which strongly affects the status of Earth’s ionosphere. Such long-term monitoring exceeds the capabilities of one single space instrument. It can only be achieved by combining observations of different missions. Together with GOES/EUVS, SOHO/SEM, TIMED/SEE, SORCE/SOLSTICE, and SDO/EVE, LYRA is part of the picture. In addition, most of its channels are interesting candidates to serve as proxies for reconstructing the whole solar irradiance in the EUV (see Dudok de Wit *et al.*, 2009 and Cessateur *et al.*, 2011).

#### 4.3. Sun–Moon Eclipses

LYRA measurements of irradiance profiles during solar eclipses are a matter of special interest. The irradiance data acquired during solar eclipses allow us to assess the center-to-limb variations (CLV) of the solar brightness. As the radiation from different disk positions originates at different heights, the CLV curves allow to sample the broad range of layers of the solar atmosphere. This provides useful information for testing and refining solar atmosphere models (*cf.* Neckel, 2005; Koesterke, Allende Prieto, and Lambert, 2008).

Up to mid-2012, LYRA observed seven solar eclipses: two in 2010 (15 January and 11 July), four in 2011 (4 January, 1 June, 1 July, and 25 November), and one in 2012 (20 May). The 15 January 2010 eclipse was the longest annular solar eclipse of this millennium. The radiative transfer COde for Solar Irradiance (COSI: Haberreiter, Schmutz, and Hubeny, 2008 and Shapiro *et al.*, 2010) was used to model the CLV of the solar radiation measured in the LYRA Herzberg channel and to model the light curves of the 15 January 2010 eclipse. The comparison of measured and calculated light curves is presented in Figure 15. One can see that the calculations with COSI are in excellent agreement with the LYRA measurements; see Shapiro *et al.* (2011) for a more complete analysis of eclipses.

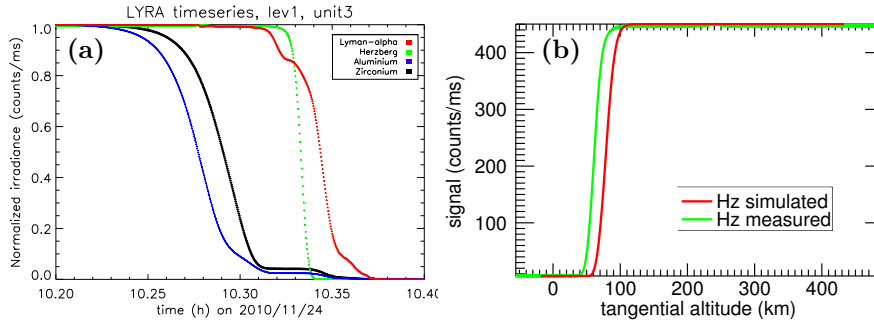


**Figure 15.** Relative irradiance data of Herzberg channel during the annular eclipse on 15 January 2010. Time is in hours after 15 January 2010 00:00 UT. (a) Passage 1 observed by unit 1 and unit 2. (b) Passage 2 observed only by unit 2. The crosses indicate the profiles calculated with the COSI code.

#### 4.4. Occultations

As the Sun approaches the Earth's shadow, LYRA's line-of-sight crosses deeper layers of the Earth's atmosphere. There, atmospheric constituents absorb part of the solar radiation in a range of wavelengths that is intrinsically dependent of their chemical nature (Figure 16 (a)). Analyzing the extinction curve of each LYRA channel therefore provides information about the distribution of some of these constituents as a function of altitude. The main molecules active in LYRA bandpasses are O, O<sub>2</sub>, O<sub>3</sub>, N<sub>2</sub>, and H<sub>2</sub>O. In Figure 16 (b), we provide an example of the reconstruction of Herzberg-signal extinction by O<sub>2</sub> and O<sub>3</sub> (see Dominique *et al.* (2009), and references therein).

Observing the extinction curves also allows one to detect the contamination of a spectral channel by longer wavelengths. Monitoring the evolution of such a curve therefore provides a reliable way to track any spectral degradation of the channel and hopefully to correct for it.



**Figure 16.** Panel (a) shows the normalized extinction curves when LYRA is progressively observing the Sun through deeper layers of the Earth’s atmosphere (*i.e.* when the tangential altitude is decreasing), for all channels. Panel (b) shows comparison between modeled and measured Herzberg-channel extinction curve for increasing tangential altitudes.

#### 4.5. Space Weather Monitor

The solar EUV flux informs us about the occurrence of solar events and forms a vital input for aeronomy models (Hochedez *et al.*, 2005). Consequences of such events might affect, among others, telecommunication or spacecraft trajectories (Bothmer and Daglis, 2007). With its quasi-continuous visibility of the Sun (the above-mentioned short occultation periods excepted), its high-cadence and its approximately nine contacts with the ground per day, LYRA constitutes a valuable space-weather monitor.

### 5. Conclusion

After more than two years of operations, LYRA demonstrated its ability to produce time series of solar irradiance in the XUV–EUV–MUV range with a very favorable sampling rate. Uncalibrated and calibrated data products, as well as various quicklooks and a list of flares, are distributed via the instrument website (<http://proba2.sidc.be>). Among others, the instrument appears to be particularly well adapted to the analysis of flares, for which its spectral channels and sampling rate are ideal. High sampling rate also benefits other fields of research, of which some have been mentioned in Section 4 (analysis of the center-to-limb brightness variation, determining the composition of Earth’s upper atmosphere). Finally, LYRA finds its place in the long-term monitoring of solar irradiance in the EUV–UV range.

**Acknowledgements** The authors acknowledge the support from the Belgian Federal Science Policy Office through the ESA-PRODEX programme. LYRA is a project of the Centre Spatial de Liège, the Physikalisch-Meteorologisches Observatorium Davos, and the Royal Observatory of Belgium, funded by the Belgian Federal Science Policy Office (BELSPO) and by the Swiss Bundesamt für Bildung und Wissenschaft. The authors want to acknowledge this

consortium, as well as the whole LYRA team for their investment and motivation. A special thanks goes to D. McMullin and J. Zender for their pertinent suggestions.

## References

- Allende Prieto, C., Asplund, M., Fabiani Bendicho, P.: 2004, Center-to-limb variation of solar line profiles as a test of NLTE line formation calculations. *Astron. Astrophys.* **423**, 1109–1117. doi:10.1051/0004-6361:20047050.
- BenMoussa, A., Schühle, U., Haenen, K., Nesládek, M., Koizumi, S., Hochedez, J.-F.: 2004, PIN diamond detector development for LYRA, the solar VUV radiometer on board PROBA II. *Physica Status Solidi App. Res.* **201**, 2536–2541. doi:10.1002/pssa.200405187.
- BenMoussa, A., Theissen, A., Scholze, F., Hochedez, J.F., Schühle, U., Schmutz, W., Haenen, K., Stockman, Y., Soltani, A., McMullin, D., Vest, R.E., Kroth, U., Laubis, C., Richter, M., Mortet, V., Gissot, S., Delouille, V., Dominique, M., Koller, S., Halain, J.P., Remes, Z., Petersen, R., D’Olieslaeger, M., Defise, J.-M.: 2006, Performance of diamond detectors for VUV applications. *Nuc. Inst. Meth. Phys. Res. A* **568**, 398–405. doi:10.1016/j.nima.2006.06.007.
- BenMoussa, A., Dammasch, I.E., Hochedez, J.-F., Schühle, U., Koller, S., Stockman, Y., Scholze, F., Richter, M., Kroth, U., Laubis, C., Dominique, M., Kretzschmar, M., Mekaoui, S., Gissot, S., Theissen, A., Giordanengo, B., Bolsee, D., Hermans, C., Gillotay, D., Defise, J.-M., Schmutz, W.: 2009, Pre-flight calibration of LYRA, the solar VUV radiometer on board PROBA2. *Astron. Astrophys.* **508**, 1085–1094. doi:10.1051/0004-6361/200913089.
- Bothmer, V., Daglis, I.A.: 2007, *Space Weather – Physics and Effects*, Praxis Publishing.
- Cessateur, G., Dudok de Wit, T., Kretzschmar, M., Lilensten, J., Hochedez, J.-F., Snow, M.: 2011, Monitoring the solar UV irradiance spectrum from the observation of a few passbands. *Astron. Astrophys.* **528**, A68. doi:10.1051/0004-6361/201015903.
- Dolla, L., Marqué, C., Seaton, D.B., Van Doorselaere, T., Dominique, M., Berghmans, D., Cabanas, C., De Groof, A., Schmutz, W., Verdini, A., West, M.J., Zender, J., Zhukov, A.: 2012, Time delays in quasi-periodic pulsations observed during the X2.2 solar flare on 15 February 2011. *Astrophys. J. Lett.* **749**, L16. doi:10.1088/2041-8205/749/1/L16.
- Dominique, M., Gillotay, D., Cessateur, G., Cabanas, C., Dammasch, I., De Groof, A., Kretzschmar, M., Nicula, B., Tetard, C., Vanhellemont, F., Zender, J.: 2009, The contribution of PROBA2-LYRA occultations to Earth atmosphere composition analysis, *TBD TBD* 285–295.
- Dudok de Wit, T., Kretzschmar, M., Lilensten, J., Woods, T.: 2009, Finding the best proxies for the solar UV irradiance. *Geophys. Res. Lett.* **361**, L10107. doi:10.1029/2009GL037825.
- Fontenla, J.M., Curdt, W., Haberreiter, M., Harder, J., Tian, H.: 2009, Semiempirical Models of the Solar Atmosphere. III. Set of Non-LTE Models for Far-Ultraviolet/Extreme-Ultraviolet Irradiance Computation. *Astrophys. J.* **707**, 482–502. doi:10.1088/0004-637X/707/1/482.
- Haberreiter, M.: 2011, Solar EUV Spectrum Calculated for Quiet Sun Conditions. *Solar Phys.*, **274**, 473–479. ADS:2011SoPh..274..473H, doi:10.1007/s11207-011-9767-9.
- Haberreiter, M., Schmutz, W., Hubeny, I.: 2008, NLTE model calculations for the solar atmosphere with an iterative treatment of opacity distribution functions. *Astron. Astrophys.* **492**, 833–840. doi:10.1051/0004-6361:200809503.
- Hochedez, J.-F., Zhukov, A., Robbrecht, E., van der Linden, R., Berghmans, D., Vanlommel, P., Theissen, A., Clette, F.: 2005, Solar weather monitoring. *Ann. Geoph.* **23**, 3149–3161. doi:10.5194/angeo-23-3149-2005.
- Hochedez, J.-F., Schmutz, W., Stockman, Y., Schühle, U., Benmoussa, A., Koller, S., Haenen, K., Berghmans, D., Defise, J.-M., Halain, J.-P., Theissen, A., Delouille, V., Slemzin, V., Gillotay, D., Fussen, D., Dominique, M., Vanhellemont, F., McMullin, D., Kretzschmar, M., Mitrofanov, A., Nicula, B., Wauters, L., Roth, H., Rozanov, E., Rüedi, I., Wehrli, C., Soltani, A., Amano, H., van der Linden, R., Zhukov, A., Clette, F., Koizumi, S., Mortet, V., Remes, Z., Petersen, R., Nesládek, M., D’Olieslaeger, M., Roggen, J., Rochus, P.: 2006, LYRA, a solar UV radiometer on PROBA2. *Adv. Space Res.* **37**, 303–312. doi:10.1016/j.asr.2005.10.041.
- Judge, D.L., McMullin, D.R., Ogawa, H.S., Hovestadt, D., Klecker, B., Hilchenbach, M., Mobius, E., Canfield, L.R., Vest, R.E., Watts, R., Tarrío, C., Kuehne, M., Wurz, P.: 1998, First Solar EUV Irradiances Obtained from SOHO by the CELIAS/SEM. *Solar Phys.* **177**, 161–173. ADS:1998SoPh..177..161J, doi:10.1023/A:1004929011427.

- Koesterke, L., Allende Prieto, C., Lambert, D.L.: 2008, Center-to-Limb Variation of Solar Three-dimensional Hydrodynamical Simulations. *Astrophys. J.* **680**, 764–773. doi:10.1086/587471.
- Krivova, N.A., Solanki, S.K., Wenzler, T., Podlipnik, B.: 2009, Reconstruction of solar UV irradiance since 1974. *J. Geophys. Res. (Atmospheres)* **114**, D13. doi:10.1029/2009JD012375.
- Lean, J., Rottman, G., Harder, J., Kopp, G.: 2005, SORCE Contributions to New Understanding of Global Change and Solar Variability. *Solar Phys.* **230**, 27–53. ADS:2005SoPh..230..27L, doi:10.1007/s11207-005-1527-2.
- Neckel, H.: 2005, Analytical Reference Functions  $F(\lambda)$  for the Sun's Limb Darkening and Its Absolute Continuum Intensities ( $\lambda\lambda$  300 to 1100 m). *Solar Phys.* **229**, 13–33. ADS:2005SoPh..229...13N, doi:10.1007/s11207-005-4081-z.
- Rottman, G.J., Woods, T.N., Sparr, T.P.: 1993, Solar-Stellar Irradiance Comparison Experiment 1. I - Instrument design and operation. *J. Geophys. Res.* **98**, 10667–10677. doi:10.1029/93JD00462.
- Schmutz, W., Fehlmann, A., Hülsen, G., Meindl, P., Winkler, R., Thuillier, G., Blattner, P., Buisson, F., Egorova, T., Finsterle, W., Fox, N., Gröbner, J., Hochedez, J.-F., Koller, S., Meftah, M., Meissonnier, M., Nyeki, S., Pfiffner, D., Roth, H., Rozanov, E., Spescha, M., Wehrli, C., Werner, L., Wyss, J.U.: 2009, The PREMOS/PICARD instrument calibration. *Metrologia* **46**, 202–206. doi:10.1088/0026-1394/46/4/S13.
- Seaton, D., Berghmans, D., Nicula, B., Halain, J.P., De Groof, A., Thibert, T., Bloomfield, D.S., Raftery, C.L., Gallagher, P.T., Auchère, F., Defise, J.-M., D'Huys, E., Lecat, J.-H., Mazy, E., Rochus, P., Rossi, L., Schühle, U., Slemzin, V., Yalim, M.S., Zender, J.: 2012, The SWAP EUV Imaging Telescope Part I: Instrument Overview and Pre-Flight Testing. *SPh* doi:10.1007/s11207-012-0114-6.
- Shapiro, A.I., Schmutz, W., Schoell, M., Haberreiter, M., Rozanov, E.: 2010, NLTE solar irradiance modeling with the COSI code. *Astron. Astrophys.* **517**, A48. doi:10.1051/0004-6361/200913987.
- Shapiro, A.I., Schmutz, W., Dominique, M., Shapiro, A.V.: 2011, Eclipses observed by LYRA - a sensitive tool to test the solar atmosphere models. *Solar Phys.* doi:10.1007/s11207-012-0063-0.
- Van Doorselaere, T., De Groof, A., Zender, J., Berghmans, D., Goossens, M.: 2011, LYRA Observations of Two Oscillation Modes in a Single Flare. *Astrophys. J.* **740**, 90. doi:10.1088/0004-637X/740/2/90.
- Woodraska, D.L., Woods, T.N., Eparvier, F.G.: 2004, In-flight calibration and performance of the Solar Extreme ultraviolet Experiment (SEE) aboard the TIMED Satellite. In: Nardell, C.A., Lucey, P. G., Yee, J.-H., Garvin, J. B. (eds.) *SPIE*, **CS-5660**, 36–47. doi:10.1117/12.579034.
- Woods, T.N., Rottman, G.J., Ucker, G.J.: 1993, Solar-Stellar Irradiance Comparison Experiment 1. II - Instrument calibrations. *J. Geophys. Res.* **98**, 10679–10694. doi:10.1029/93JD00463.
- Woods, T., Eparvier, F., Hock, R., Jones, A., Woodraska, D., Judge, D., Didkovsky, L., Lean, J., Mariska, J., Warren, H., McMullin, D., Chamberlin, P., Berthiaume, G., Bailey, S., Fuller-Rowell, T., Sojka, J., Tobiska, W., Viereck, R.: 2012, Extreme Ultraviolet Variability Experiment (EVE) on the Solar Dynamics Observatory (SDO): Overview of Science Objectives, Instrument Design, Data Products, and Model Developments. *Solar Phys.* **275**, 115–143. ADS:2012SoPh..275..115W, doi:10.1007/s11207-009-9487-6.

